

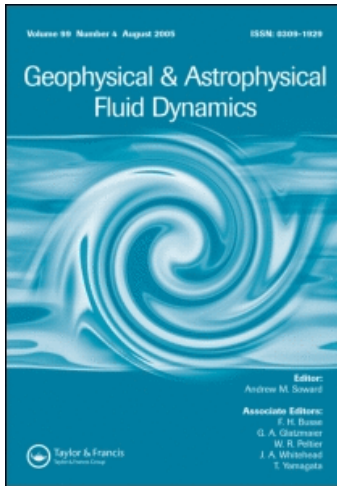
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Geophysical & Astrophysical Fluid Dynamics

Publication details, including instructions for authors and subscription information:

<http://www.informaworld.com/smpp/title~content=t713642804>

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Piyali Chatterjee^a; Axel Brandenburg^{ab}; Gustavo Guerrero^a

^a NORDITA, AlbaNova University Center, SE 10691 Stockholm, Sweden ^b Department of Astronomy, AlbaNova University Center, Stockholm University, SE 10691 Stockholm, Sweden

First published on: 13 August 2010

To cite this Article Chatterjee, Piyali, Brandenburg, Axel and Guerrero, Gustavo(2010) 'Can catastrophic quenching be alleviated by separating shear and α effect?', *Geophysical & Astrophysical Fluid Dynamics*, 104: 5, 591 – 599, First published on: 13 August 2010 (iFirst)

To link to this Article: DOI: 10.1080/03091929.2010.504185

URL: <http://dx.doi.org/10.1080/03091929.2010.504185>

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Can catastrophic quenching be alleviated by separating shear and α effect?

PIYALI CHATTERJEE*[†], AXEL BRANDENBURG^{†‡}
and GUSTAVO GUERRERO[†]

[†]NORDITA, AlbaNova University Center, Roslagstullsbacken 23,
SE 10691 Stockholm, Sweden

[‡]Department of Astronomy, AlbaNova University Center, Stockholm University,
SE 10691 Stockholm, Sweden

(Received 14 April 2010; in final form 14 June 2010)

The small-scale magnetic helicity produced as a by-product of the large-scale dynamo is believed to play a major role in dynamo saturation. In a mean-field model the generation of small-scale magnetic helicity can be modelled by using the dynamical quenching formalism. Catastrophic quenching refers to a decrease of the saturation field strength with increasing Reynolds number. It has been suggested that catastrophic quenching only affects the region of non-zero helical turbulence (i.e. where the kinematic α operates) and that it is possible to alleviate catastrophic quenching by separating the region of strong shear from the α layer. We perform a systematic study of a simple axisymmetric two-layer $\alpha\Omega$ dynamo in a spherical shell for Reynolds numbers in the range $1 \leq R_m \leq 10^5$. In the framework of dynamical quenching we show that this may not be the case, suggesting that magnetic helicity fluxes would be necessary.

Keywords: Dynamo; Magnetic helicity; Catastrophic quenching

1. Introduction

It is widely believed that the solar magnetic cycle is driven by an $\alpha\Omega$ dynamo. The region of strong radial shear called the tachocline (Spiegel and Zahn 1992) at the bottom of the solar convection zone (SCZ) is believed to be the place where strong toroidal field is formed due to stretching of the weaker but diffuse poloidal field. It has been inferred from helioseismology that the tachocline is confined within a thin layer in the overshoot region which lies below the SCZ. This led Parker (1993) to propose the idea of an interface dynamo, where the shear is confined to a region with a greatly reduced turbulent diffusion, which is also the region of production of strong toroidal field. The helical turbulence generated due to convection and rotation in the layer above provides the turbulent α and a large turbulent diffusivity η_t . The dynamo cycle being thus completed, the interface dynamo operates as a surface wave propagating along the

*Corresponding author. Email: piyalic@nordita.org

boundary between strong shear and convection, which is also a region with a strong gradient in the turbulent diffusivity.

In order to mimic the process of dynamo saturation, traditionally authors have used an algebraic quenching function like $\alpha_0/(1 + \overline{\mathbf{B}}^2/B_{\text{eq}}^2)$ or even $\alpha_0/(1 + R_m \overline{\mathbf{B}}^2/B_{\text{eq}}^2)$, where α_0 is the unquenched value of α , R_m is the magnetic Reynolds number, $\overline{\mathbf{B}}$ is the mean magnetic field and B_{eq} is the equipartition magnetic field. However, it appears that conservation of magnetic helicity plays an important role in the process of saturation. At large R_m , the magnetic helicity ($\int \mathbf{A} \cdot \mathbf{B} dV$) is fairly well conserved by the dynamo producing equal amounts of helicity in small and large scales, respectively. If the small-scale magnetic helicity is not able to escape out of the system, then the turbulent α effect is markedly reduced (Pouquet *et al.* 1976). This leads to the magnetic energy of the dynamo to be quenched such that the saturation value varies as R_m^{-1} . Since astrophysical objects have large R_m ($R_m \sim 10^9$ for the Sun), this strong dependence of the saturation field strength B_{sat} on R_m is referred to as *catastrophic quenching*.

Even though the helicity constraint in direct numerical simulations (DNS) of dynamos with strong shear has been clearly identified, the results can be matched with mean-field models having a weaker algebraic quenching of α and turbulent diffusivity than α^2 dynamos (Brandenburg *et al.* 2001). However, the empirically determined coefficients would depend on circumstances and are therefore not universal. Such a model would not obey magnetic helicity evolution and is therefore untenable on theoretical grounds. The interface dynamo model has been invoked several times as a way of getting around the persistent problem of catastrophic quenching (see Charbonneau 2005 for a review). The belief is that the quenching function remains close to unity in the region of finite α since the toroidal field is expected to be weak there. However, to our knowledge, this has never been verified in a consistent manner for a range of magnetic Reynolds numbers. In this article we perform a series of calculations with mean-field $\alpha\Omega$ models, in spherical geometry, considering both algebraic and dynamical quenching formulations, for magnetic Reynolds numbers in the range $1 \leq R_m \leq 2 \times 10^5$. An important feature of these models is that the region of strong narrow shear is separated from the region of helical turbulence, as proposed in the Parker's interface model.

In section 2 we discuss the features of the $\alpha\Omega$ model used, and the formulation of dynamical α quenching. The results are highlighted in section 3 and conclusions are drawn in section 4. A part of the calculations presented in this article will be discussed in a more detailed article (Chatterjee *et al.* 2010). In this article we focus specifically on the reality of the catastrophic quenching in dynamos with α and Ω effects operating in two widely separated layers.

2. Nonlinear $\alpha\Omega$ dynamo

2.1. The underlying mean-field model

Our dynamo model consists of the induction equations for the toroidal component of the mean poloidal field potential, $A_\phi(r, \theta)$, and the mean toroidal magnetic field, $B_\phi(r, \theta)$, written in spherical geometry under the assumption of axisymmetry ($\partial/\partial\phi = 0$). In some of the cases an additional evolution equation will be solved for the α effect

(described in the next section). We have used a modified version of the publicly available solar dynamo code *Surya* (the code *Surya* and its manual can be obtained by writing an email to arnab@physics.iisc.ernet.in) described in Chatterjee *et al.* (2004) to perform these calculations.

In this article we have used a smoothed step profile for η given by

$$\eta(r) = \eta_r + \frac{1}{2}\eta_t \left[1 + \operatorname{erf}\left(\frac{r-r_e}{d_e}\right) \right], \quad (1)$$

where $r_e = 0.73R_\odot$ and $d_e = 0.025R_\odot$. We define the magnetic Reynolds number as $R_m = \eta_t/\eta_r$. In order to facilitate comparison with earlier work in Cartesian geometry (see, e.g. Brandenburg *et al.* 2009), it is convenient to define an effective minimal wavenumber, k_1 . Somewhat arbitrarily, we use $k_1 = 2/R_\odot$, which corresponds to a harmonic wave with 2 nodes spanning the full latitudinal extent between both poles. We also define the wavenumber of the energy-carrying eddies, k_f , corresponding to the inverse pressure scale height near the base of the convection zone. For all our calculations we have taken $k_f = 7k_1$. Using the estimate $\eta_t = u_{\text{rms}}/3k_f$ we can express the equipartition field strength with respect to the turbulent kinetic energy as

$$B_{\text{eq}} = (4\pi\rho)^{1/2}u_{\text{rms}} = (4\pi\rho)^{1/2}3\eta_t k_f.$$

For algebraic quenching there is no magnetic α effect and we just have the kinetic α effect, α_K , which we assume to be of the form

$$\alpha_K(r) = \frac{1}{2}\alpha_0 \cos\theta \left[1 + \operatorname{erf}\left(\frac{r-r_a}{d_a}\right) \right] / \left(1 + g_\alpha \bar{B}^2/B_{\text{eq}}^2 \right), \quad (2)$$

where the value of α_0 may be computed using the first-order smoothing approximation (FOSA) as being equal to $\varepsilon_f \eta_t k_f$ (Blackman and Brandenburg 2002). Here, the prefactor ε_f is usually of order 0.1 or less since $(\mathbf{u} \cdot \boldsymbol{\omega})_{\text{rms}} < u_{\text{rms}}\omega_{\text{rms}}$. The case $\varepsilon_f = 1$ means the flow is maximally helical. The term g_α is a non-dimensional coefficient equal to 1 or R_m depending on the assumed form of algebraic quenching in the models. Even though the helical turbulence pervades almost the entire convection zone, we take $r_a = 0.77R_\odot$ and $d_a = 0.015R_\odot$ so that we can have a large separation between the shear layer and the layer where turbulence is important. Consequently we consider a differential rotation profile like that in the high latitude tachocline of the Sun given by,

$$\Omega(r) = -\frac{1}{2}\Omega_0 \left[1 + \operatorname{erf}\left(\frac{r-r_w}{d_w}\right) \right], \quad (3)$$

where $\Omega_0 = 14\text{nHz}$, $r_w = 0.68R_\odot$ and $d_w = 0.015R_\odot$. The radial profiles of η_t , α and $\partial\Omega/\partial r$ are plotted as a function of fractional radius r/R_\odot in figure 1. The region of strong radial shear is thus separated from the region of helical turbulence and the diffusivity has a strong gradient at a radius lying between the two layers. It may be noted that, in order to have supercritical dynamo action in a model with η having the same radial profile as α , we must set $\varepsilon_f \gg 1$. If the strong gradient of η lies between the two source regions, then we can work with $\varepsilon_f \leq 1$. Also the time period T_{cyl} of the oscillatory dynamo remains a reasonably small fraction of the turbulent diffusion time t_{diff} . We can justify the profiles of η and α on the grounds of having dynamo action for a reasonable range of parameters even while avoiding any significant overlap between

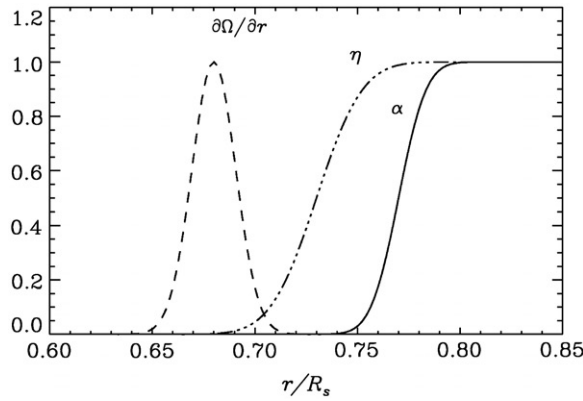


Figure 1. Profiles of negative radial shear $\partial\Omega/\partial r$ (dashed), α (solid) and η (dashed-dotted) as a function of fractional solar radius.

the two source regions. The formulation of the equation for the evolution of the α effect for dynamical quenching is described in section 2.2.

2.2. Dynamical α quenching

It was first shown by Pouquet *et al.* (1976) that the turbulent α effect is modified due to the generation of small-scale magnetic helicity such that the total α effect is given by,

$$\alpha = \alpha_K + \alpha_M = -\frac{\tau}{3}(\overline{\boldsymbol{\omega} \cdot \mathbf{u}} - \rho^{-1} \overline{\mathbf{j} \cdot \mathbf{b}}), \quad (4)$$

where $\boldsymbol{\omega}$, \mathbf{u} , \mathbf{j} , and \mathbf{b} denote the fluctuating components of vorticity, velocity, current density, and magnetic field in the plasma, respectively.

For this type of quenching we use the same radial and latitudinal profiles of α_K as given in equation (2), but without the algebraic quenching factor in the denominator, i.e. we put $g_\alpha = 0$. The second term in equation (4) is sometimes referred to as the magnetic α effect or α_M . It is possible to write an equation for the evolution of α_M from the equation for the evolution of the small-scale magnetic helicity density $h_f = \overline{\mathbf{a} \cdot \mathbf{b}}$ using the relation (Brandenburg *et al.* 2009),

$$\alpha_M = \frac{\eta_t k_f^2}{B_{\text{eq}}^2} h_f. \quad (5)$$

The equation for $\overline{\mathbf{a} \cdot \mathbf{b}}$ is in principle gauge-dependent. However, under the assumption of scale separation, i.e. when the correlation length of the turbulence is small compared to the system size, one can define a magnetic helicity density of small-scale fields in a gauge-independent manner as the density of linkages (Subramanian and Brandenburg 2006). Using equation (5), this leads to an evolution equation for α_M ,

$$\frac{\partial \alpha_M}{\partial t} = -2\eta_t k_f^2 \left(\frac{\overline{\boldsymbol{\mathcal{E}} \cdot \overline{\mathbf{B}}} + \frac{\alpha_M}{R_m}}{B_{\text{eq}}^2} \right) - \nabla \cdot \mathbf{F}_\alpha, \quad (6)$$

where $\bar{\mathcal{E}}$ and $\bar{\mathbf{B}}$ are the mean electromotive force and the mean magnetic field. The flux \mathbf{F}_α consists of individual contributions, e.g. advection due to the mean flow, the Vishniac–Cho flux (Vishniac and Cho 2001, Subramanian and Brandenburg 2004), diffusive fluxes, triple correlation terms, etc. The effect of some of these individual fluxes in a spherical geometry and with both radial and latitudinal shear have been investigated by Guerrero *et al.* (2010), but for all the simulations presented in this article we have put $\mathbf{F}_\alpha = 0$.

We solve the equations for $A_\phi(r, \theta)$, $B_\phi(r, \theta)$ and α_M in a domain confined by $0 \leq \theta \leq \pi$ and $0.55R_\odot \leq r \leq R_\odot$. The boundary conditions for A_ϕ are given by a potential field condition at the surface (Dikpati and Choudhuri 1994) and $A_\phi = 0$ at the poles. At the bottom we use the perfect conductor boundary condition $\partial(rB_\theta)/\partial r = \partial(rB_\phi)/\partial r = 0$. Also $B_\phi = 0$ on all other boundaries. We have checked that the results are not very sensitive to different boundary conditions at the bottom boundary mainly because the boundary is far removed from the dynamo region. Since $\mathbf{F}_\alpha = 0$, no derivatives of α_M need to be evaluated, so no boundary conditions need to be specified for α_M and its evolution equation is just an initial value problem. We start with an initial dipolar solution where B_ϕ is antisymmetric about the equator.

3. Results

To study the R_m dependence of B_{sat} in our model we keep all the dynamo parameters the same for all the runs except η_r , which we change from $2 \times 10^5 \text{ cm}^2 \text{ s}^{-1}$ to $2 \times 10^{10} \text{ cm}^2 \text{ s}^{-1}$ while keeping η_t fixed at $4 \times 10^{10} \text{ cm}^2 \text{ s}^{-1}$. It may also be noted that the time period of the dynamo models (T_{cyl}) is fairly independent of the magnetic Reynolds number.

To be able to correctly compare the dynamo models for different R_m , we have calculated the critical value of α_0 , denoted by α_c for each model. In the following we present results for $\alpha_0 = 2\alpha_c$. We show in figure 2 the butterfly diagrams for B_ϕ and α_M at a depth of $0.72R_\odot$. It may be concluded from the butterfly diagram for α_M in figure 2(b) that the small-scale current helicity, and hence α_M , is predominantly negative (positive) in the Northern (Southern) hemisphere. Let us denote the exponential decay time for α_M by t_α . So, $t_\alpha = R_m/\eta_t k_f^2 = 4.55 \times 10^{-3} R_m t_{\text{diff}}$. For $R_m = 2 \times 10^3$, the decay time $t_\alpha \gg T_{\text{cyl}}$ and so the system of equations is underdamped, as can be seen from the butterfly diagrams in figure 2(a) as well as from saturation curve (dashed dotted line) in figure 3. Note that there are amplitude modulations of the magnetic field before it settles to a final saturation value. The nature of the saturation curves of the magnetic energy is thus strongly governed by the ratio of t_α and T_{cyl} .

The results of our calculations for different Reynolds numbers are plotted in figure 3. The slopes in the kinematic phase are similar for all R_m within the error in the numerical determination of the critical α_c . The strong R_m dependence, which is reminiscent of catastrophic quenching in large R_m dynamos, can be easily discerned from the same figure, but is more clear in figure 4 where we see that the saturation energy decreases monotonically as a function of magnetic Reynolds number. For $R_m = 2 \times 10^5$, the code has to be run for 500 t_{diff} before the dynamo fields may start becoming “strong” again for the case with $\alpha_0 = 2\alpha_c$. Due to long computational times involved in this exercise we have not continued the calculation beyond 60 t_{diff} . Hence the determination of saturation magnetic energy may be inaccurate for $R_m = 2 \times 10^5$. In figure 4 we compare

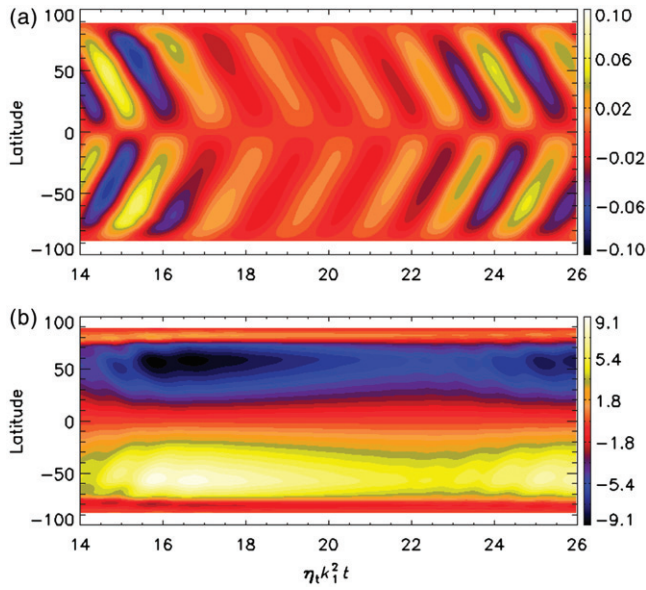


Figure 2. (a) $B_\phi(0.72R_\odot, \theta)$ and (b) $\alpha_m(0.72R_\odot, \theta)$ as a function of diffusion time $\eta_t k_1^2 t$ for $R_m = 2 \times 10^3$.

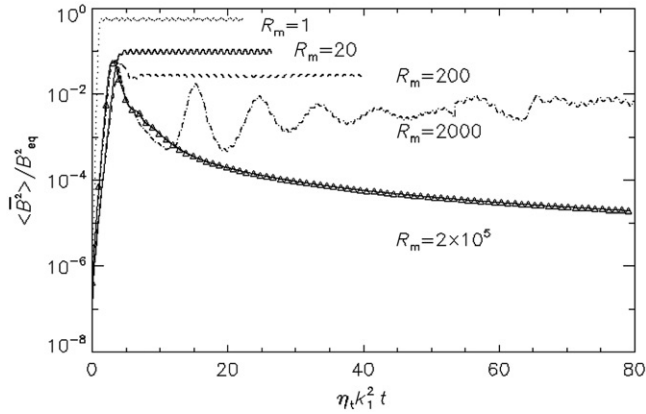


Figure 3. Volume averaged magnetic energy in the domain scaled with the equipartition energy for $R_m = 1$ (dotted line), $R_m = 20$ (solid), $R_m = 200$ (dashed), $R_m = 2 \times 10^3$ and $R_m = 2 \times 10^5$ (triangles) with dynamical quenching.

the case with dynamical quenching against the cases with a simple algebraic quenching of the form given in equation (2) with $g_\alpha = 1$ and $g_\alpha = R_m$. We notice that for $g_\alpha = R_m$, the algebraically and dynamically quenched α effects seem to give similar dependences on R_m .

When the code is run longer, we start seeing changes in the parity after $t > 40 t_{\text{diff}}$ for the dynamically quenched system in contrast to the “strong” algebraic quenching case, where the parity remains dipolar. However, the magnetic energy and the dynamo period T_{cyl} remain fairly constant even while the system fluctuates between symmetric and anti-symmetric parity at an irregular time interval.

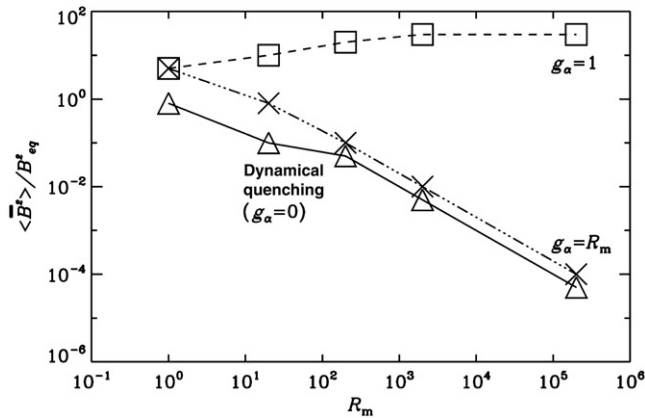


Figure 4. Volume averaged magnetic energy scaled with the equipartition energy in the saturation phase as a function of R_m for dynamical α quenching (triangles + solid) and algebraic quenching with $g_\alpha = 1$ (squares + dashed) and with $g_\alpha = R_m$ (cross + dashed-dotted). Adapted from Chatterjee *et al.* (2010).

In figure 5 we show the meridional snapshots in the Northern hemisphere of the toroidal component of the magnetic field and α_M . It may be noted that the regions strongest in B_ϕ become progressively confined in the narrow shear layer with increasing R_m while α_M becomes stronger leading to decrease in B_{sat} . Even though α_M is predominantly negative in the Northern hemisphere there is a region of positive small-scale helicity generated just below the region where α_K is finite so that contribution from the term αB_ϕ^2 in the source $\vec{\mathcal{E}} \cdot \vec{\mathbf{B}}$ is small. This effect is similar to the one reported in Brandenburg *et al.* (2009). Also we have not observed any evidence of chaotic behaviour in the range of magnetic Reynolds number $20 \leq R_m \leq 2 \times 10^5$ for twice supercritical α ($\alpha \leq 2\alpha_c$), in agreement with Covas *et al.* (1997).

4. Conclusions

The calculations done in this article indicate that it is not possible to escape catastrophic quenching due to accumulation of small-scale helicity in the domain by merely separating the regions of shear and α effect. The saturation value of magnetic energy decreases as $\sim R_m^{-1}$ for both dynamical quenching and the “strong” ($\sim R_m$ -dependent) form of algebraic quenching for the simple two-layer model. However, additionally we observe parity fluctuations for cases with dynamical quenching. It does not seem to us that there exists any chaotic behaviour in the time series of magnetic energy since the dynamo period and the saturation energy remains fairly constant. It may be possible that solar wind, coronal mass ejections, and Vishniac and Cho fluxes help in throwing out the small scale helicity from the Sun and thus alleviate catastrophic quenching. In this study we have not found any difference between the nature of the saturation curves for an $\alpha\Omega$ dynamo and an α^2 dynamo using the form of dynamical quenching given by equation (6). However, it is clear that the algebraic quenching formula must fail if we were to allow for magnetic helicity fluxes that would, under suitable circumstances, alleviate catastrophic quenching.

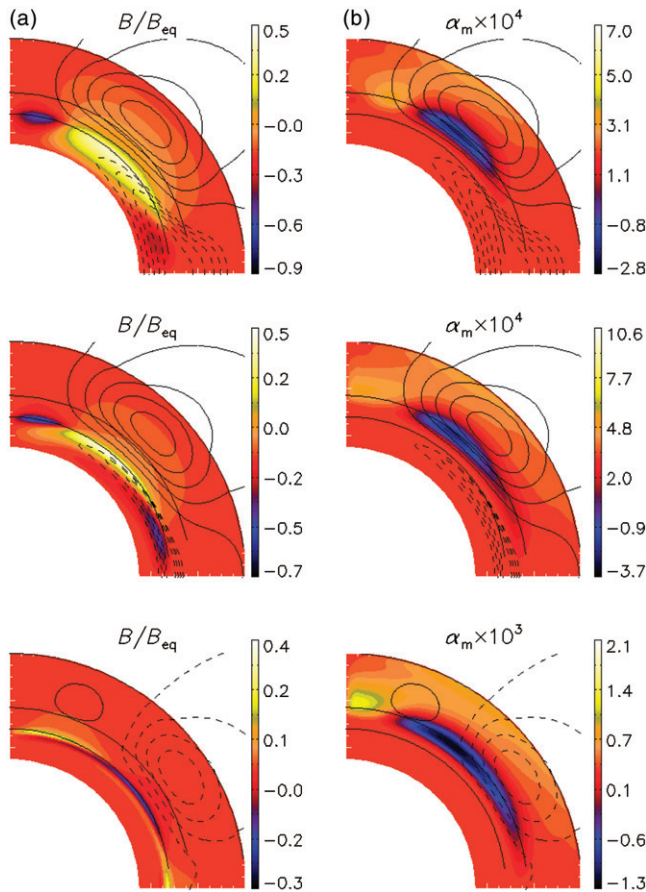


Figure 5. Meridional snapshots of (a) B_ϕ/B_{eq} and (b) α_m (color/grey-scale coded) for cases with different R_m starting from 20 (upper panel), 200 (middle panel) and 2×10^3 (lower panel). Two concentric circles have been drawn at $0.68R_\odot$ and $0.77R_\odot$ to denote the radial positions of the shear layer and the α effect. Solid and dashed lines denote poloidal field lines, corresponding to contours of positive (negative) $r \sin \theta A_\phi$.

We have been cautious about using dynamical quenching equation for dynamo numbers not very large compared to the critical dynamo numbers. We would expect that the magnetic field should affect all the turbulence coefficients including both α and η . However for this analysis we have not included any quenching of η_t . This may be justified since such an effect could be mimicked by our simple two-layer model with a lower η_t in the region of production of strong toroidal fields and a higher η_t in the region of weaker poloidal fields. The effect of dynamical quenching on more realistic solar dynamo models having meridional circulation, Babcock–Leighton α effect, and diffusive helicity fluxes have also been studied (Chatterjee *et al.* 2010).

Unfortunately the direct numerical simulations have not yet reached the modest Reynolds numbers used in this article ($\sim 10^4$) which are still much lower than the astrophysical dynamos. If B_{sat}^2 really has an inverse dependence on R_m , then the solar dynamo should not be operating like it does! We may conclude that either we must have helicity fluxes out of the system or cross-equatorial diffusive fluxes inside the domain or

that the α_M equation to be used beside the mean-field induction equations must be suitably modified for $\alpha\Omega$ dynamos. To verify if the equation for dynamical quenching works in the same way as in α^2 dynamos, we need to perform systematic comparisons between DNS with shear and convection and mean-field modelling for $\alpha\Omega$ dynamos.

Acknowledgements

P.C. and A.B. would like to thank the organisers of the “Natural Dynamos” meeting in High Tatras, where this work was initiated. We acknowledge the allocation of computing resources provided by the Swedish National Allocations Committee at the Center for Parallel Computers at the Royal Institute of Technology in Stockholm and the National Supercomputer Centers in Linköping as well as the Norwegian National Allocations Committee at the Bergen Center for Computational Science. This work was supported in part by the European Research Council under the AstroDyn Research Project No. 227952 and the Swedish Research Council Grant No. 621-2007-4064.

References

- Blackman, E.G. and Brandenburg, A., Dynamic nonlinearity in large scale dynamos with shear. *Astrophys. J.* 2002, **579**, 359–373.
- Brandenburg, A., Bigazzi, A. and Subramanian, K., The helicity constraint in turbulent dynamos with shear. *Monthly Notices Roy. Astron. Soc.* 2001, **325**, 685–692.
- Brandenburg, A., Candelaresi, S. and Chatterjee, P., Small-scale magnetic helicity losses from a mean-field dynamo. *Monthly Notices Roy. Astron. Soc.* 2009, **398**, 1414–1422.
- Charbonneau, P., Dynamo Models of the Solar Cycle. *Living Rev. Solar Phys.* 2005, **2**, 2.
- Chatterjee, P., Guerrero, G. and Brandenburg, A., Magnetic helicity fluxes in interface and flux transport dynamos. *Astron. Astrophys.* 2010, submitted, arXiv:1005.5335.
- Chatterjee, P., Nandy, D. and Choudhuri, A.R., Full-sphere simulations of a circulation-dominated solar dynamo: Exploring the parity issue. *Astron. Astrophys.* 2004, **427**, 1019–1030.
- Covas, E., Tworkowski, A., Brandenburg, A. and Tavakol, R., Dynamos with different formulations of a dynamic α -effect. *Astron. Astrophys.* 1997, **317**, 610–617.
- Dikpati, M. and Choudhuri, A.R., The evolution of the Sun’s poloidal field. *Astron. Astrophys.* 1994, **291**, 975–989.
- Guerrero, G., Chatterjee, P. and Brandenburg, A., Shear-driven and diffusive helicity fluxes in $\alpha\Omega$ dynamos. *Monthly Notices Roy. Astron. Soc.* 2010, in press, arXiv:1005.4818.
- Parker, E.N., A solar dynamo surface wave at the interface between convection and nonuniform rotation. *Astrophys. J.* 1993, **408**, 707–719.
- Pouquet, A., Frisch, U. and Léorat, J., Strong MHD helical turbulence and the nonlinear dynamo effect. *J. Fluid Mech.* 1976, **77**, 321–354.
- Subramanian, K. and Brandenburg, A., Nonlinear current helicity fluxes in turbulent dynamos and alpha quenching. *Phys. Rev. Lett.* 2004, **93**, 205001.
- Subramanian, K. and Brandenburg, A., Magnetic helicity density and its flux in weakly inhomogeneous turbulence. *Astrophys. J.* 2006, **648**, L71–L74.
- Spiegel, E.A. and Zahn, J.-P., The solar tachocline. *Astron. Astrophys.* 1992, **265**, 106–114.
- Vishniac, E.T. and Cho, J., Magnetic helicity conservation and astrophysical dynamos. *Astrophys. J.* 2001, **550**, 752–760.